

Orbital and physical parameters of eclipsing binaries from the All-Sky Automated Survey catalogue

VI. AK Fornacis – a rare, bright K-type eclipsing binary[★]

K. G. Hełminiak^{1,2,3}, R. Brahm¹, M. Ratajczak², N. Espinoza¹, A. Jordán^{4,1}, M. Konacki^{2,5}, and M. Rabus^{1,6}

¹ Pontificia Universidad Católica de Chile, Av. Vicuña Mackenna 4860, 7820436 Macul, Santiago, Chile
e-mail: xysiek, rbrahm, nespino, ajordan, mrabus@astro.puc.cl

² Nicolaus Copernicus Astronomical Center, Department of Astrophysics, ul. Rabiańska 8, 87-100 Toruń, Poland
e-mail: xysiek, milena, maciej@ncac.torun.pl

³ Subaru Telescope, National Astronomical Observatory of Japan, 650 North Aohoku Place, Hilo, HI 96720, USA

⁴ Millennium Institute of Astrophysics, Av. Vicuña Mackenna 4860, 7820436 Macul, Santiago, Chile

⁵ Astronomical Observatory, A. Mickiewicz University, ul. Słoneczna 36, 60-286 Poznań, Poland

⁶ Max-Planck-Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany

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ABSTRACT

Aims. We present the results of the combined photometric and spectroscopic analysis of a bright ($V = 9.14$), nearby ($d = 31$ pc), late-type detached eclipsing binary AK Fornacis. This $P = 3.981$ d system has not been previously recognised as a double-lined spectroscopic binary, and this is the first full physical model of this unique target.

Methods. With the FEROS, CORALIE and HARPS spectrographs we collected a number of high-resolution spectra in order to calculate radial velocities of both components of the binary. Measurements were done with our own disentangling procedure and the TODCOR technique, and were later combined with the photometry from the ASAS and SuperWASP archives. We also performed an atmospheric analysis of the component spectra with the Spectroscopy Made Easy (SME) package.

Results. Our analysis shows that AK For consists of two active, cool dwarfs having masses of $M_1 = 0.6958 \pm 0.0010$ and $M_2 = 0.6355 \pm 0.0007 M_\odot$ and radii of $R_1 = 0.687 \pm 0.020$ and $R_2 = 0.609 \pm 0.016 R_\odot$, slightly less metal abundant than the Sun. Parameters of both components are well reproduced by the models.

Conclusions. AK For is the brightest system among the known eclipsing binaries with K or M type stars. Its orbital period is one of the longest and rotational velocities one of the lowest, which allows us to obtain very precise radial velocity measurements. The precision in physical parameters we obtained places AK For among the binaries with the best mass measurements in the literature. It also fills the gap in our knowledge of stars in the range of $0.5\text{--}0.8 M_\odot$, and between short and long-period systems. All this makes AK For a unique benchmark for understanding the properties of low-mass stars.

Key words. binaries: eclipsing – binaries: spectroscopic – stars: fundamental parameters – stars: late-type – stars: individual: AK Fornacis

1. Introduction

During the last years we have witnessed a remarkable improvement in the study of lower main-sequence stars, mainly by the increasing number of eclipsing binaries found to contain K and M type dwarfs. The interest in those objects has increased recently with the new exoplanet search surveys dedicated especially to them, such as the WFCAM Transit Survey¹. The knowledge of the host star parameters, especially mass and radius, is necessary to constrain the characteristics of the planet (Demory et al. 2011). Another justification for studying late type stars are the still unexplained discrepancies between the models and observations. We still do not fully understand why in most cases, sometimes surprisingly, the stars are larger and cooler than predicted (e.g. Lacy 1977; Popper 1997; Torres & Ribas 2002; Morales

et al. 2009; Irwin et al. 2011), while in others the same characteristics are nicely reproduced (Thompson et al. 2010; Feiden et al. 2011; Hełminiak & Konacki 2011). Unfortunately, most of newly discovered systems tend to be faint, and have large rotational velocities due to short periods ($P < 3$ d), thus their properties are difficult to measure accurately. Only a handful of them have their masses and radii determined with the accuracy better than 3% (with systematics taken into account properly), thus allowing for meaningful tests of the stellar models (Torres et al. 2010). Each new system is highly-valuable, and in this paper we present the analysis of the brightest one known to date, which also has one of the most accurate mass measurements among the eclipsing binaries until now.

The variable nature of AK Fornacis (HD 21703, HIP 16247, ASAS J032923-2406.1; hereafter AK For) was discovered by the *Hipparcos* mission, and the star was classified as a probably eclipsing in the 74th Special Name-list of Variable Stars (Kazarovets et al. 1999). Before that it was only known as a K-type star (Upgren et al. 1972) and an X-ray source (Gioia et

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¹ <http://www.ast.cam.ac.uk/~sth/wts/index.html>

Table 1. Literature and catalogue data for AK For

Parameter	Value	Ref.
α_{ICRS}	03:29:22.87422	1
δ_{ICRS}	-24:06:03.0926	1
π [mas]	32.19(1.13)	1
μ_α [mas/yr]	220.58(95)	1
μ_δ [mas/yr]	99.53(1.66)	1
B [mag]	10.46	2
V_{Mer} [mag]	9.36	2
V_{ASAS} [mag]	9.135	3
J [mag]	7.031(21)	4
H [mag]	6.510(53)	4
K [mag]	6.262(33)	4
Sp. Type	K3 V k Fe+0.4	5
P_{orb} [d]	3.981	6

Notes. 1: *Hipparcos* (van Leeuwen 2007); 2: Mermilliod (1986); 3: our fit to ASAS data; 4: 2MASS (Skrutskie et al. 2006); 5: Gray et al. (2006); 6: Otero (2003).

al. 1990). Otero (2003) gave the first period estimation – 3.981 d – on the basis of *Hipparcos* and the All-Sky Automated Survey (ASAS; Pojmański 2002) data available that time, however AK For is not listed in the ASAS Catalogue of Variable Stars. The only spectroscopic study was done in low resolution by Gray et al. (2006), who classified AK For as a K3 V, “chromospherically very active” star. Finally, Bailer-Jones (2011) estimated the effective temperature and extinction towards the star on the basis of multi-band photometry, but his method is not suitable for binaries, so both values are overestimated (5350 K, 1.28 mag). The summary of catalogue and literature data is presented in Table 1. Note the discrepancy between ASAS and Mermilliod’s (1986) values of the V magnitude. The system’s T_{eff} calculated from $V - K$ colours and calibrations by Worthey & Lee (2011) is ~ 4100 K for Mermilliod (1986) and ~ 4350 K for ASAS.

2. Observations

2.1. Spectroscopy

With the Fiber-fed Extended Range Optical Spectrograph (FEROS; Kaufer et al. 1999) we collected twelve high-resolution ($R \sim 40\,000$) spectra in November 2011, September 2012, and March 2013. They were supplemented with six $R \sim 70\,000$ spectra taken in February 2013 with the CORALIE spectrograph, attached to the 1.2-m Euler telescope in La Silla, and two $R \sim 115\,000$ spectra from the High Accuracy Radial velocity Planet Searcher (HARPS; Mayor et al. 2003) obtained in quadrature on September 09, 2012. FEROS and CORALIE data were reduced with a dedicated Python-based pipeline (for a description see: Jordán et al. 2014), which was initially built for CORALIE data reduction, but was modified to deal with FEROS data. HARPS spectra were reduced on-site with the available Data Reduction Software (DRS).

2.2. Photometry

The V -band photometry of AK For, publicly available from the ASAS Catalogue², spans from November 2000 to December 2009, and contains 883 good quality points (flagged “A” in the original data). The time span of the public Wide-Angle Search

for Planets (SuperWASP; Pollacco et al. 2006) data³ is much shorter – from August 2006 till February 2010, but with only a few, practically useless points after Feb. 2008. Due to the orbital period very close to 4 days, the eclipses are not always observable in the night, so in the case of SuperWASP we initially selected only 2007–2008 data, where eclipses were visible. We later limited ourselves to 2709 data points with individual errors smaller than 0.03 mag. This threshold level was a result of the further analysis, in which we wanted to keep as many good quality data points as possible (especially in eclipses), while at the same time minimizing the systematics in the resulting radii. We also noted that depths of single minima vary slightly, but it was impossible to conclude if it is a real effect (due to spots for example) or an uncorrected systematic. AK For is known to be an active system but no consistent out-of-eclipse modulation was found in the SuperWASP photometry.

To check for the spot-originated brightness variations, we performed V -band observations with the PISCO telescope, attached to the main tube of the 1.2-m Euler telescope in La Silla. PISCO is a 31.75 cm (12.5 in) aperture reflector by RC Optical Systems, with a Finger Lakes Instrumentation (FLI) Proline-09000 camera, equipped with a 3056×3056 pixel Kodak KAF-09000 CCD. With the pixel scale of $0.862''/\text{pix}$, the field of view is $44'' \times 44''$, strongly vignetted in the corners. It contains a set of standard Johnson-Cousins U, B, V, R, I filters, supplemented by near-IR Z and Y . It works together with the CORALIE spectrograph allowing for a simultaneous photometry and spectroscopy, taking exposures while CORALIE’s shutter is open. With PISCO we secured 58 images in the V band, with exposure times of 18 sec, taken exactly in the same time as CORALIE spectra from nights February 24, 25 and 26, 2013. Standard CCD reduction steps, utilizing bias and sky-flat images, were performed in IRAF, as well as a simple aperture photometry under the task `apphot`. AK For was by far the brightest star in the field, so the selection of comparison stars was difficult. We inspected seven other stars, and found two that did not show any significant brightness variations but unfortunately are about 2 mag fainter than AK For. We only wanted to perform a relative photometry, in order to detect any variability, so we did not transform PISCO magnitudes to the standard Johnson’s system, nor we observed photometric standard stars.

3. Analysis

3.1. Radial velocities and the orbital solution

Radial velocities (RVs) were initially calculated with our implementation of the two dimensional cross-correlation technique (TODCOR; Zucker & Mazeh 1994) with synthetic spectra used as templates. Due to the poor S/N in the blue part, limits implemented to the CORALIE and FEROS reduction pipelines (optimized for high-precision RVs), and limitations of the synthetic templates, we used the following wavelength ranges: 4135–6500 Å for FEROS, 4400–6500 Å for CORALIE, and 3800–6500 Å for HARPS. One of the TODCOR’s features is that it also gives the most probable flux ratio of the two components. We found it to be 0.423 and that it is pretty stable along the V and SuperWASP bands. Later we used the FEROS spectra to perform disentangling, as described in Konacki et al. (2010), and obtain separate spectra of the two components. We used them to perform the spectral analysis, and for final RV measurements. For

³ <http://exoplanetarchive.ipac.caltech.edu/applications/ExoTables/search.html?dataset=superwasptimeseries>

² <http://www.astrouw.edu.pl/asas/?page=aasc>

Table 2. Individual RV measurements, errors and residuals (all in km/s) of the components of AK For, together with the orbital phase, exposure times (in seconds) and S/N of the spectra. 2F marks the MPG 2.2-m/FEROS, 3H the ESO 3.6-m/HARPS, and EC the 1.2-m Euler/CORALIE data (with “p” meaning simultaneous observations with PISCO).

JD-2450000	v_1	σ_1	$(O - C)_1$	v_2	σ_2	$(O - C)_2$	Phase	T_{exp}	S/N	Tel./Sp.
5876.681435	-37.483	0.062	-0.003	46.441	0.202	-0.008	0.0965	450	90	2F
5876.841349	-50.637	0.049	0.031	60.928	0.155	0.026	0.1366	900	125	2F
5877.644573	-57.370	0.067	-0.149	67.959	0.242	-0.122	0.3384	750	120	2F
5877.701489	-53.645	0.061	0.003	64.000	0.231	-0.165	0.3527	1200	90	2F
5877.783808	-47.668	0.038	0.018	57.725	0.139	0.092	0.3734	1200	95	2F
5878.673272	42.913	0.056	-0.001	-41.535	0.185	0.043	0.5968	600	95	2F
5878.759453	50.399	0.043	0.010	-49.798	0.126	-0.041	0.6184	450	90	2F
5878.810493	54.428	0.039	0.021	-54.178	0.137	-0.024	0.6313	480	105	2F
6179.782690	-67.426	0.095	0.004	79.214	0.329	-0.056	0.2336	900	85	3H
6179.879929	-67.687	0.089	0.027	79.484	0.339	-0.097	0.2580	780	90	3H
6195.817064	-67.545	0.048	0.080	79.397	0.134	-0.086	0.2613	480	120	2F
6345.559577	52.316	0.054	0.094	-51.594	0.136	0.075	0.8757	420	35	EC
6346.511053	-44.029	0.184	-0.132	53.513	0.294	-0.061	0.1147	600	25	EC
6347.545693	-47.323	0.060	-0.001	57.305	0.144	-0.021	0.3746	600	30	EC
6348.542794	52.439	0.060	-0.046	-52.171	0.123	-0.212	0.6251	480	30	ECp
6349.537454	52.613	0.091	0.146	-51.882	0.134	0.055	0.8749	420	25	ECp
6350.536805	-47.586	0.046	-0.069	57.613	0.093	0.072	0.1259	600	30	ECp
6376.509645	59.529	0.075	-0.198	-59.889	0.191	0.086	0.6502	600	90	2F
6377.503283	44.095	0.062	-0.054	-42.758	0.173	0.172	0.8997	600	135	2F
6383.474455	-38.866	0.068	-0.001	47.980	0.279	0.012	0.3997	600	75	2F

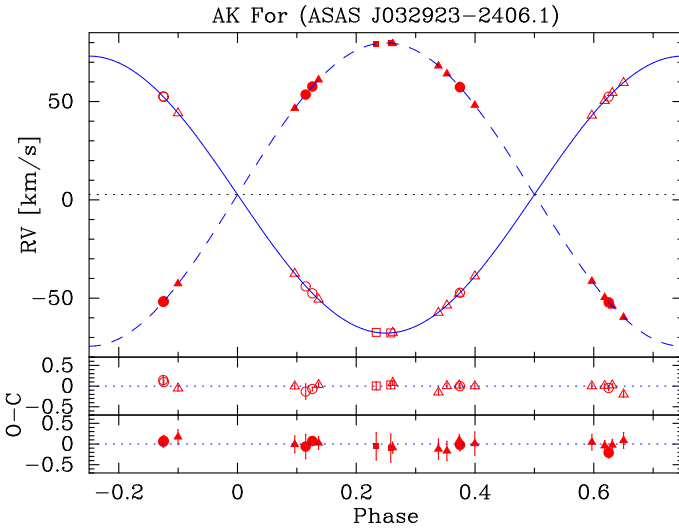


Fig. 1. Radial velocities and orbital solution for AK For. Solid line and empty symbols refer to the primary, dashed line and filled symbols to the secondary component. Triangles mark the FEROS, circles mark the CORALIE and squares the HARPS data. The systemic velocity is marked with the dotted line. The residuals are plotted below together with the (scaled) individual errors. The *rms* of the orbital fit is 82 and 96 m/s for the primary and secondary respectively.

the FEROS data we followed the procedure of least-squares fitting described in Konacki et al. (2010), and for CORALIE and HARPS we used the disentangled spectra as templates in TODCOR.

The RV measurements were analysed with a simple procedure, which fits a double-keplerian orbit using the Levenberg-Marquardt algorithm. As free parameters we set the orbital period P , the time of primary conjunction T_0 , corresponding in this case to the primary (deeper) eclipse, the velocity semi-amplitudes $K_{1,2}$, primary’s systemic velocity γ_1 , the difference in components systemic velocities $\gamma_2 - \gamma_1$, and the difference

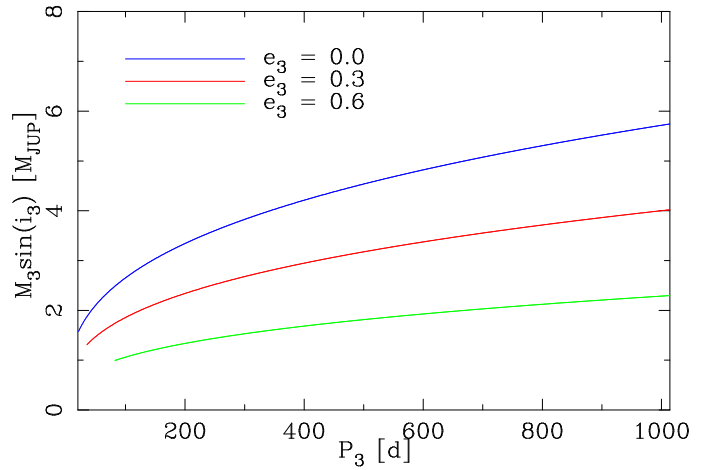


Fig. 2. Third body detection $M_3 \sin(i_3)$ limits estimated from our RV measurements, for several values of outer orbit’s eccentricities. We should be able to detect bodies with masses above these limits.

between FEROS and CORALIE zero-points 2F–EC, measured separately for each component. We assumed that HARPS has the same zero-point as FEROS, as we did not have enough HARPS data to securely constrain this difference. HARPS measurements do not outlay significantly from the final solution, so we find this assumption correct. Initially, we have also fitted for the eccentricity e and the argument of the periastron ω but neither in this step, nor in any other, we found the orbit eccentric. Thus we kept $e = 0$ during the whole process. Finally, we ended up with 40 data points (20 for each component) and 8 parameters.

To estimate reliable parameter uncertainties, with systematics accounted for, we run 10000 bootstrap iterations. All the RV measurements, together with their uncertainties, residuals from the model RV curve, orbital phases, exposure times and S/N around 5500Å, are shown in the Table 2. The resulting orbital parameters are presented in Table 3.

The precision of RVs we have reached is good enough to look for massive circumbinary planets around the system. If we assume that we can detect any RV variation of the semi-amplitude of at least 96 m/s – the *rms* of the secondary – we can calculate the mas of a body that could produce such a variation. We show those calculations in Figure 2 as a function of the outer orbit’s period and for three values of the outer orbit’s eccentricity: 0.0, 0.3 and 0.6. The value of the 96 m/s was assumed to be the RV semi-amplitude. The lower limit of the period is 21 d (the major semi-axis of 34.83 R_{\odot}) which refers to the shortest stable circular orbit (Holman & Wiegert 1999). For eccentric orbits the limits are terminated at the shortest periods having the distance of periastron larger than 34.83 R_{\odot} (around 35 and 82 d). The upper limit of the period is 1014 d, twice the time span of our data. One can see that our radial velocities are precise enough to detect massive planets on stable orbits around AK For. If there was a third body orbiting within this range of periods, it must have the mass lower than the limit.

3.2. Atmospheric parameters

We performed a spectral atmospheric analysis using the Spectroscopy Made Easy package (hereafter SME; Valenti & Piskunov 1996). The disentangled spectra were first scaled with the flux ratio obtained from TODCOR (0.423). We’ve also run SME with other flux ratios, but the resulting temperatures led to further inconsistencies, in distances for example.

Following the work of Valenti & Fischer (2005), in order to obtain a continuum placement and derived parameters more accurate, we have chosen seven FEROS orders between 5317 and 6397 Å, each of them was analysed separately. We used the list of atomic lines from Vienna Atomic Line Database (VALD; Piskunov et al. 1995; Kupka et al. 1999) generated for the Sun for the initial values described by Valenti & Fischer (2005). We also adopted atmosphere models of Kurucz (1993). We set the $\log(g)$ for each component to the values obtained with JKTEBOP (see: Tab. 3 and next Section), micro- and macro-turbulence velocities to 1 and 3 km/s respectively, while T_{eff} , $[M/H]$ and $v \sin(i)$ were fitted for every order. As starting values of temperatures we used 4500 and 4150 K for the primary and secondary respectively. For each order we also started with three different values of $[M/H]$: -1, -0.5 and 0. From all SME runs we took only those, for which the resulting $v \sin(i)$ was close to the value obtained from JKTEBOP (see next Section). Of the 21 SME runs for each component, only 13 and 9 was successful for the primary and secondary respectively. Example of a successful fit to one of the orders of the primary’s spectrum is presented in Figure 3.

For the adopted values of effective temperatures, metallicities and projected rotational velocities we took the averages, and standard deviations as their uncertainties. Those values are given in Table 3. The ratio of velocities is in agreement with the ratio of radii found in the light curve fitting (see next Section), but the velocities themselves are larger than predicted by the final solution and tidal locking. This could have been caused by underestimation of the assumed micro- or macro-turbulence velocity. What is surprising we found two significantly different values of $[M/H]$: -0.3 and -0.1 for the primary and secondary respectively, both with uncertainties of 0.05 dex. We found that this difference was not dependent on the starting values of any parameter, including the flux ratio, and was not correlated with the resulting temperatures. We need to note however, that the FEROS spectra were not corrected for the scattered light, which might have influenced the SME analysis, and that the S/N of the

Table 3. Orbital and physical parameters of AK For.

AK For Parameter	Primary		Secondary	
	Value	\pm	Value	\pm
<i>Radial velocity fit</i>				
P [d]	3.9809913 ± 0.0000045			
T_0 [JD]	2451903.2681 ± 0.0045			
K [km/s]	70.471	0.031	77.136	0.051
γ [km/s]	2.667	0.028	2.516	0.115
$M \sin^3(i)$ [M_{\odot}]	0.6937	0.0010	0.6336	0.0008
2F–EC [km/s]	-0.037	0.032	0.133	0.073
2F–3H [km/s]	0.0	—	0.0	—
rms [km/s]	0.082		0.096	
<i>SME spectral analysis</i>				
T_{eff} [K]	4690	100	4390	150
$[M/H]$	-0.29	0.05	-0.10	-0.05
$v \sin(i)$ [km/s]	9.52	0.22	8.27	0.42
<i>Light curve fit^a</i>				
i [°]	87.428 ± 0.080			
r [R/a]	0.0591	0.0017	0.0524	0.0014
J_{ASAS}	0.528 ± 0.028			
J_{SW}	0.592 ± 0.030			
<i>Adopted physical parameters</i>				
a [R_{\odot}]	11.6318 ± 0.0047			
M [M_{\odot}]	0.6958	0.0010	0.6355	0.0007
R [R_{\odot}]	0.687	0.020	0.609	0.016
$\log g$	4.607	0.025	4.673	0.023
v_{rot} [km/s] ^b	8.72	0.25	7.73	0.21
$\log(L/L_{\odot})$	-0.687	0.045	-0.908	0.064
M_{bol} [mag]	6.47	0.11	7.02	0.16
<i>From and based on Hipparcos</i>				
d [pc]	31.1 ± 1.1			
M_V [mag]	7.06	0.08	7.91	0.09
<i>From JKTEBSDIM</i>				
d [pc]	32.4 ± 1.6			
M_V [mag]	6.91	0.17	7.68	0.27

Notes. ^(a) Combined ASAS and SuperWASP solutions; ^(b) assuming synchronization;

secondary’s spectrum is significantly lower than the primary’s. We thus adopt a conservative value of $[M/H] = -0.2 \pm 0.1$ dex.

3.3. Light curve solution

For the light curve (LC) analysis we used the latest version (v28) of the code JKTEBOP (Southworth et al. 2004a,b), which is based on the EBOP program (Popper & Etzel 1981). On the basis of spectroscopic data we first found the mass ratio and ephemeris, which we included in the LC analysis. We found that the orbital period found directly by JKTEBOP from the ASAS photometry is in agreement with the one from RVs, however with larger uncertainties, and leading to significantly worse orbital solution. The SuperWASP data constrain the period even worse. For JKTEBOP we also used flux ratios found in TODCOR, and the logarithmic limb darkening (LD) law with coefficients interpolated from the tables of van Hamme (1996) for ASAS and Pollacco et al. (2006) for SuperWASP. The gravity darkening coefficients and bolometric albedos were always kept fixed at the values appropriate for stars with convective envelopes ($g = 0.32$, $A = 0.5$). As mentioned before, various tests performed on every data set did not show a significant eccentricity of the orbit of AK For, thus e was kept fixed to 0 in the

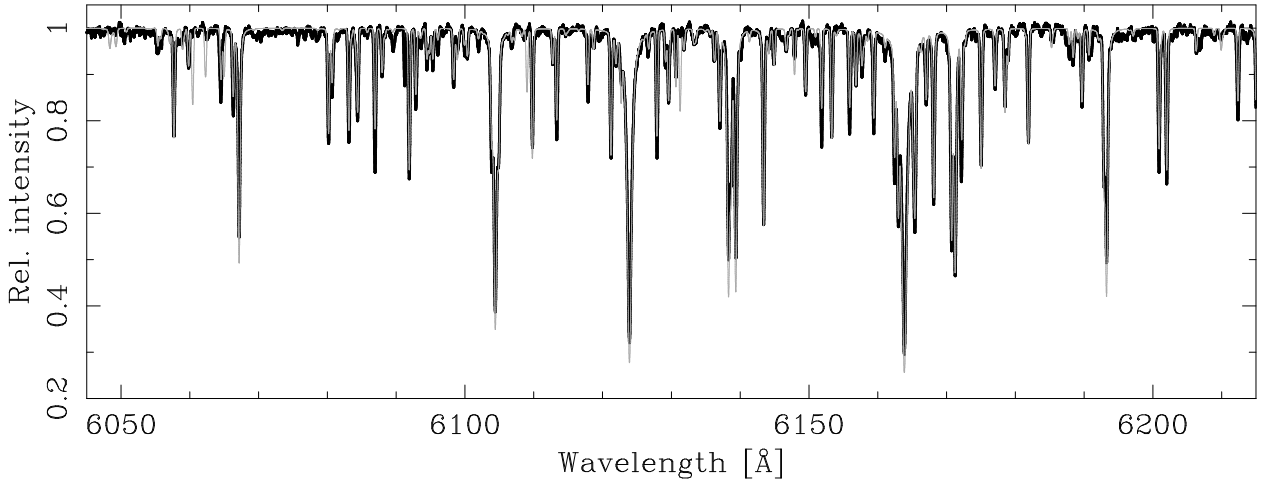


Fig. 3. Portion of the SME best fit (thin grey line) to the disentangled spectrum of the primary (thick black line).

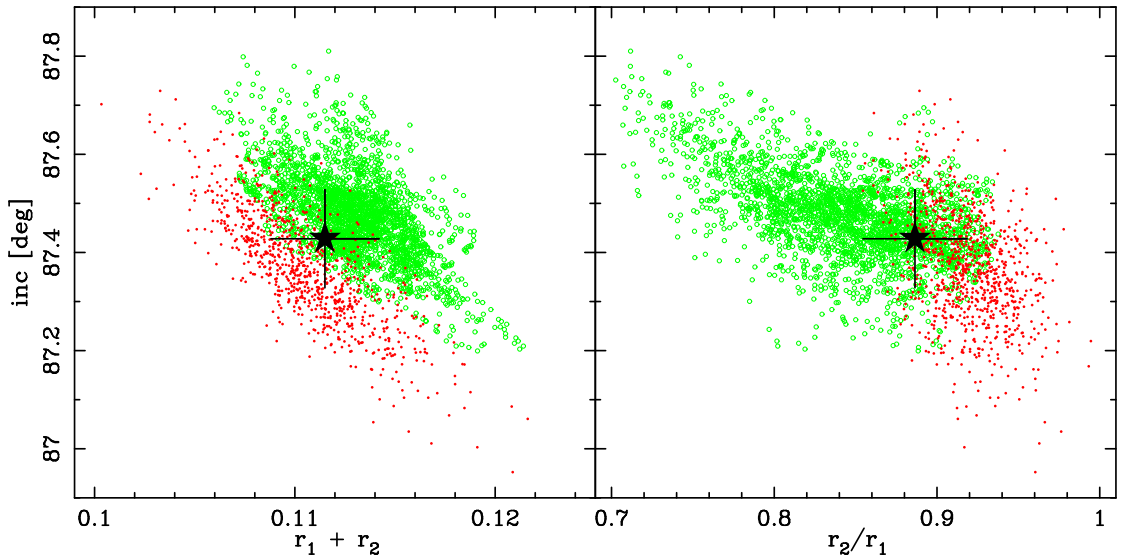


Fig. 4. Results of the residual-shifts analysis performed with JKTEBOP on the ASAS (red) and SuperWASP data (green). Plots present the distribution of consecutive solutions on the $r_1 + r_2$ vs. i (left) and $k = r_2/r_1$ vs. i (right) panels. Black stars with error bars correspond to the adopted values with their 1σ uncertainties.

analysis. We fitted for the sum of the fractional radii $r_1 + r_2$, their ratio k , orbital inclination i , surface brightness ratios J , and brightness scales (out-of-eclipse magnitudes in each filter).

We run the task 9, which to calculate reliable errors uses the residual-shifts method (Southworth 2008) to assess the importance of the correlated “red” noise, especially strong in the SuperWASP data (Southworth et al. 2011). We have run several tests to check how the final model varies with various LD coefficients and ephemeris, but we did not notice a strong dependence. Other sources of errors contribute more, but to at least partially account for LD coefficients and ephemeris uncertainties, we let them to be perturbed in the residual-shifts simulations. It is a known fact that for systems with partial eclipses, as AK For, the orbital inclination is correlated with the radii-related parameters, especially their sum. In Figure 4 we show the results of the JKTEBOP analysis on the $r_1 + r_2$ vs. i , and $k = r_2/r_1$ vs. i diagrams. We see that different data sets give similar values of inclination, but different radii, nevertheless in agreement with each other. The probable reason for this slight inconsistency is the fact that we used the same flux ratio for the two such different filters. However, as mentioned before, no significant change

of the flux ratio was noticed in the wavelengths corresponding to the SuperWASP filter. These differences can also result from the activity and the location of spots (not included here).

As the result we adopted weighted averages of the values found from the two data sets. We mark them in Figure 4, together with the adopted 1σ errors. The model LCs for ASAS and SuperWASP photometry are presented in Fig. 5. The resulting values of fractional radii $r_{1,2}$, the inclination and surface brightness ratios are given in Table 3.

The absolute values of parameters were calculated with the JKTEBOP procedure, available together with JKTEBOP. We assumed $E(B - V) = 0.0$, JHK photometry from 2MASS (Skrutskie et al. 2006), and $[M/H] = 0.0$ or -0.5 – values closest to our results of the spectral analysis. In both cases the results were practically undistinguishable. Distances were calculated with various bolometric corrections for various filters (Bessell et al. 1998; Flower 1996; Girardi et al. 2002) and surface brightness- T_{eff} relations from Kervella et al. (2004). As the result we adopted their average and the standard deviation as the uncertainty. The value of 32.4 ± 1.6 pc is in agreement with the distance from *Hipparcos* (31.1 ± 1.1 pc; van Leeuwen 2007),

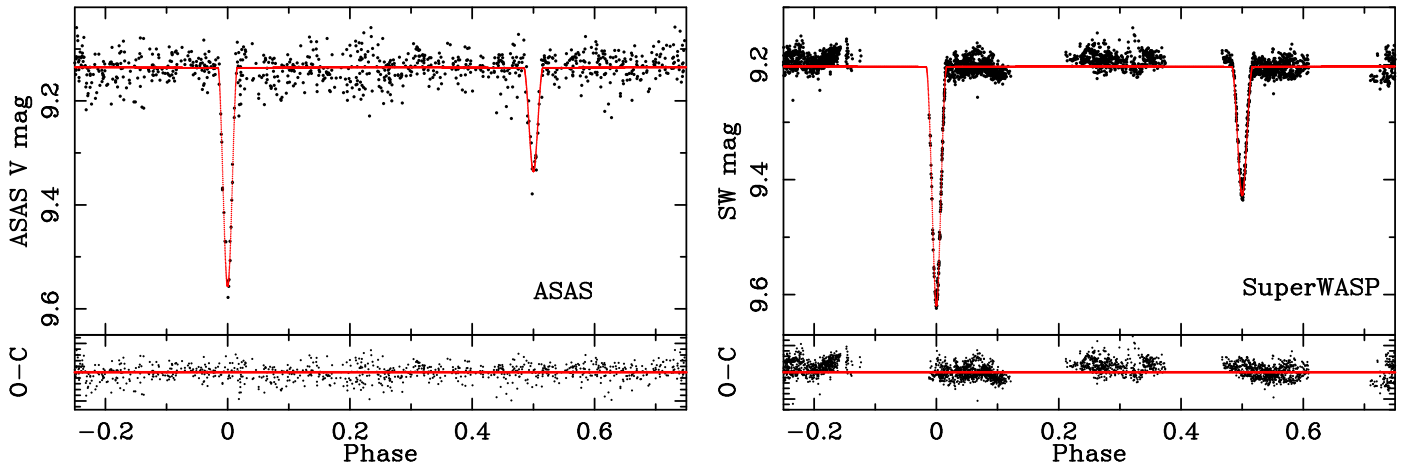


Fig. 5. Phase-folded light curves of AK For from ASAS (left) and SuperWASP (right) together with the best-fitting model and residuals.

which proves that the flux ratio from TODCOR and the temperatures we obtained in SME are correct. The single absolute V magnitudes are also in agreement with values calculated on the basis of the *Hipparcos* parallax and flux ratio from TODCOR, however the latter have smaller errors. The complete set of final parameters is presented in the Table 3. For the consistency we show distance and absolute V magnitudes obtained in two ways, but adopt the *Hipparcos*-based values as the final ones. In this way we have a complete set of absolute physical parameters of the system not based on any models or calibrations.

4. Discussion

4.1. Activity

The activity of AK For is undoubted, and was already noted (Gray et al. 2006). The system has been detected in the X-rays by the *Einstein* mission (Gioia et al. 1990). We found the emission in H_α (Fig. 6) and, particularly strong, in the Ca II H and K lines (Fig. 7; covered in our data only by HARPS). Both H_α lines have an interesting double-peak profile, with core absorption and emission in wings. The Ca II lines of the primary appear more prominent, but taking into account the flux ratio of two components around 3950 Å, it is the secondary's lines that are stronger with respect to the component's continuum. The secondary also has stronger H_α , but the H_α emission varies in the time scale of single days for both components (Fig. 6).

The long-cadence ASAS photometry (Fig. 8) shows a clear fading trend by about 0.02 mag until JD~2453500, probably originating in the evolution of spots. The additional short-cadence photometry we made with PISCO (Fig. 9) also proves the presence of spots, causing brightness variations of a similar scale as the long-term trend seen in the ASAS data. It is impossible however to constrain the spot pattern, or even to point out the component on which they are present – most likely on both. If AK For's brightness is still at the level from September 2009 (the end of ASAS data), both components are probably substantially and uniformly covered with spots. Additional, continuous, high-precision photometry from a global telescope network is required to constrain the pattern and location of spots in a given time. With such data it would be also possible to significantly improve the precision in radii, and make AK For one of the best-studied eclipsing systems.

4.2. Kinematics

The main sequence evolutionary stage of AK For is suggested by its galactic kinematics. We used our determinations of the systemic velocity together with the position, proper motion and distance from *Hipparcos* (Tab. 1). The obtained values of $U = 17.8 \pm 0.9$ km/s, $V = -7.5 \pm 0.7$ km/s and $W = 25.4 \pm 0.7$ km/s put marginally AK For in the galactic thin disk (Seabroke & Gilmore 2007). This suggests the age below 4.5-5 Gyr, however older stars are observed in the thin disk (Allende Prieto 2010; and references therein). Taking into account the large value of W , at the edge of the thin disk distribution, a young age (below 1 Gyr) seems to be unlikely.

4.3. Comparison with theoretical models

We reached a very high precision of $0.14 \pm 0.11\%$ in the masses and also a good precision in radii ($2.9 \pm 2.7\%$), allowing for meaningful tests of the stellar evolution models. AK For has the most precise mass measurements among low-mass eclipsing binaries, and it is the brightest one ($V = 9.135$ mag from the ASAS light curve). It is one of the most accurately measured eclipsing binaries in general. The precision in masses was possible due to relatively low rotational velocities, related to the period longer than for the majority of similar DEBs. Improving the accuracy of the radii and related radiative parameters will require more precise and better-calibrated photometry.

In Fig. 10 we compare our measurements with the Dartmouth stellar evolution models (Dotter et al. 2007) in the mass M vs. R , $\log(T_{\text{eff}})$, $\log(L/L_\odot)$ and M_V planes. We show 1 and 6 Gyr isochrones for three values of $[Fe/H]$ spanning from -0.1 to -0.3 dex⁴.

One can see that the 6 Gyr isochrone for the $[Fe/H]$ value obtained in the SME analysis fits to all the parameters within 1σ . The system still resides on the main sequence, but has evolved and the 1 Gyr models clearly do not fit the data. This shows how useful is the spectral analysis in solving the age-metallicity degeneracy on the main sequence. The age of 6 Gyr was found to be the best-fitting one, especially on the M/R plane. It is in agreement with the kinematic analysis described in the previous section.

⁴ SME gives $[M/H]$ but in case of the Dartmouth set the transformation to $[Fe/H]$ gives $-0.194^{+0.097}_{-0.099}$, thus we adopted $[Fe/H] = -0.2 \pm 0.1$

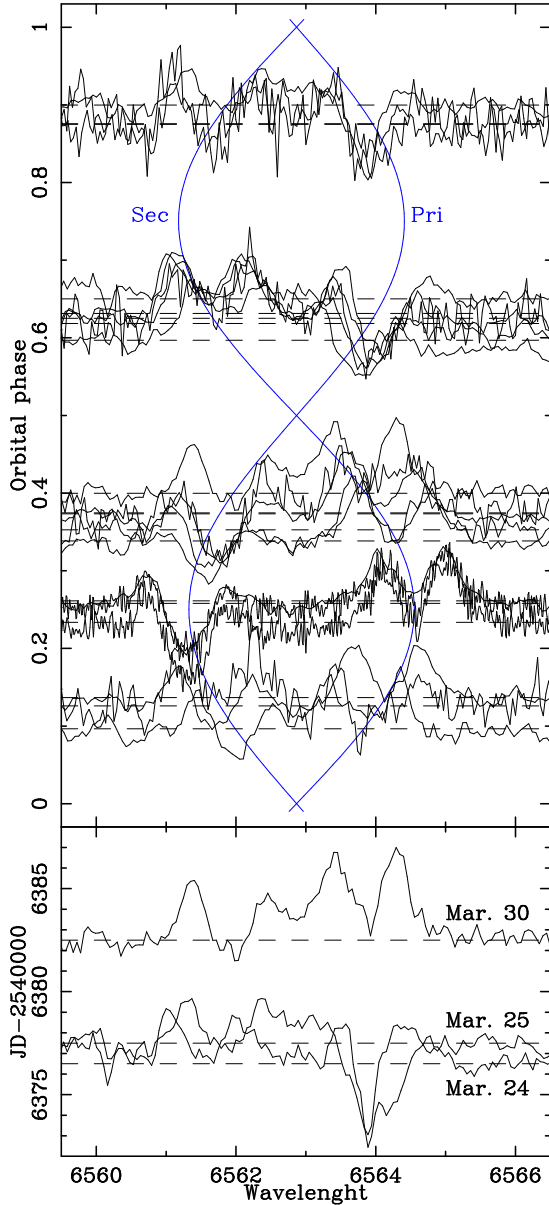


Fig. 6. H_α emission lines of AK For. In a given panel, all spectra were continuum-fitted and equally scaled. The dashed lines show the level of continuum and mark the orbital phase or time of the observation. *Top:* All spectra are shown as a function of the orbital phase. Blue solid lines trace the position of the H_α line as predicted by the orbital solution. *Bottom:* Three FEROS spectra from March 2013, showing rapid evolution of the H_α profiles. In March 30 the emission was much stronger than few days earlier. Secondary’s emission is the stronger one.

4.4. $0.4 - 0.9 M_\odot$ stars in DEBs

AK Fornacis is an unique system for several reasons. First, it is bright and nearby, allowing for further detailed studies, which may however be problematic due to the unfortunate orbital period. Second, its components masses fall in the relatively poorly studied range of $0.5-0.8 M_\odot$ – between usually very active and mostly convective M-dwarfs, and mostly inactive solar analogues of the G type. The activity is thought to be the major factor responsible for the observed discrepancies between measurements and theory in binaries, which tend to be larger for later spectral types. It is notable that in this case the theory predicts the measurements reasonably well. Finally, the orbital period of

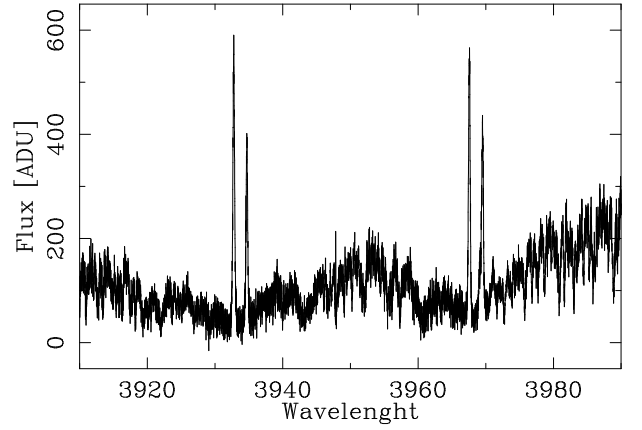


Fig. 7. Portion of HARPS spectrum of AK For around Ca II H and K lines. Emission from both components in the cores is clearly visible. The apparently stronger lines come from the primary however, referred to the fainter continuum of the secondary star, its emission lines are intrinsically the strongest. Spectrum taken in quadrature.

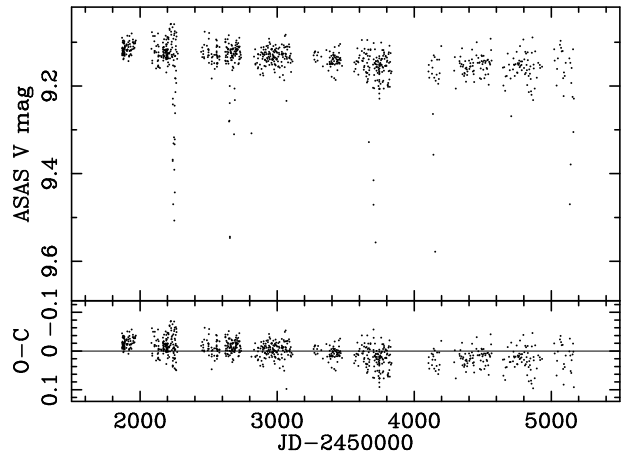


Fig. 8. Long cadence ASAS light curve (top) and the residuals of the best-fitting model from the Figure 5 (bottom) as a function of time. The change of the total brightness of the system is clearly visible, and can be explained by evolution of the spot pattern.

AK For is longer than for the majority of well-measured DEBs with K, M and late-G type components.

In the online DEBCat catalogue⁵ The notable exceptions are the secondary of M55-V54 ($0.56 M_\odot$, 9.27 d; Kałużny et al. 2014), the secondary of KIC 6131659 ($0.69 M_\odot$, 17.53 d; Bass et al. 2012), both components of LSPM J1112+7626 (much lower masses $0.39+0.27 M_\odot$, 41.03 d; Irwin et al. 2011) or both components of T-Lyr1-17236 ($0.68+0.53 M_\odot$, 8.43 d; Devor et al. 2008; not included due to poor precision in radii). It is obviously an observational selection effect and it is desirable to discover more long-period (> 3 d) low-mass eclipsing binaries.

Fig. 11 shows the mass-radius diagram for the well-measured eclipsing binaries with at least one component in the $0.4 - 0.9 M_\odot$ mass range. We show systems from the DEBCat and three examples from our previous research (Helminiak & Konacki 2011; Helminiak et al. 2011). Dartmouth isochrones for 1, 5, 10, 14 Gyr and two metallicities are plotted over. It is clearly seen that age and metallicity plays an important role

⁵ <http://www.astro.keele.ac.uk/~jkt/debcats/> of well-measured eclipsing binaries, there are almost no stars of masses lower than $0.7 M_\odot$ for periods longer than 3 days.

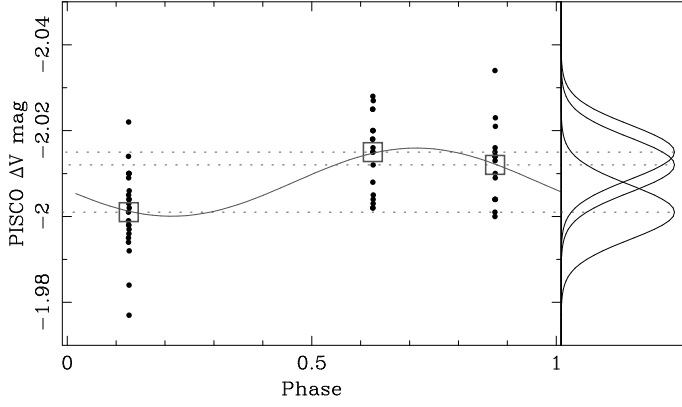


Fig. 9. PISCO V-band photometry of AK For, taken simultaneously with CORALIE observations in February 2013, phase-folded with the orbital period. For each night squares mark the average values, their error is 2 mmag (size of symbols) and the dispersion is 9 mmag, corresponding gaussians are shown to the right. The presence of spots can be deduced from a small but clear fading around phase 0.126, when presumably the spot is well visible. The brightness drop is 0.011 mag relatively to the previous night (phase 0.875). The solid line is the sine fitted to data, plotted to guide the eye only and with no physical meaning.

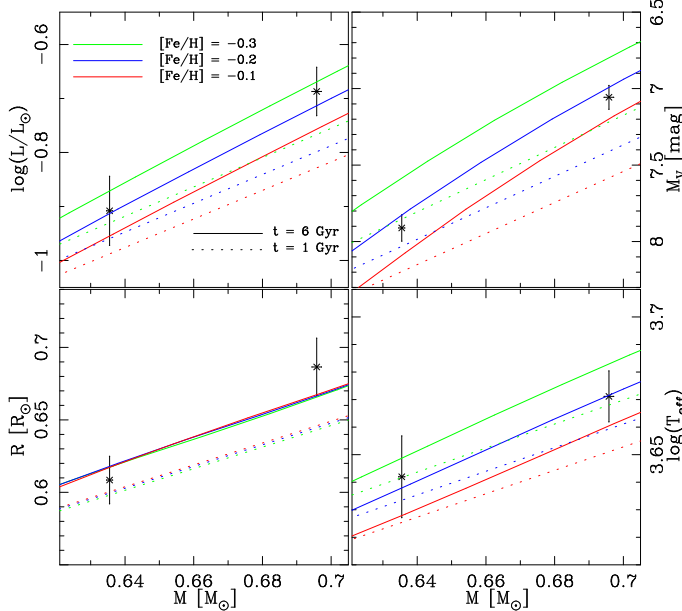


Fig. 10. Comparison of our results with the Dartmouth stellar evolution models for three values of $[Fe/H]$, and ages of 1 and 6 Gyr.

in the $M - R$ distribution, at least for stars more massive than $0.6 M_{\odot}$. Traditionally, the radii and temperatures of low mass stars were compared to the 1 Gyr, solar composition isochrone. The general characteristic was that the observed radii are larger and effective temperatures lower than models predicted. We now know a number of late type systems old enough to have their radii enlarged by the evolution (AK For should be considered one of them). In some cases metallicity can also be estimated from the spectra, which helps to constrain the age. As one can see, the traditional approach is not always valid. In order to indicate if a given star's radius is enlarged or not, one needs an independent metallicity and age estimation (possible only in clusters), or may check if the slope that the two components define on the $M - R$ diagram is compatible with any normal evolu-

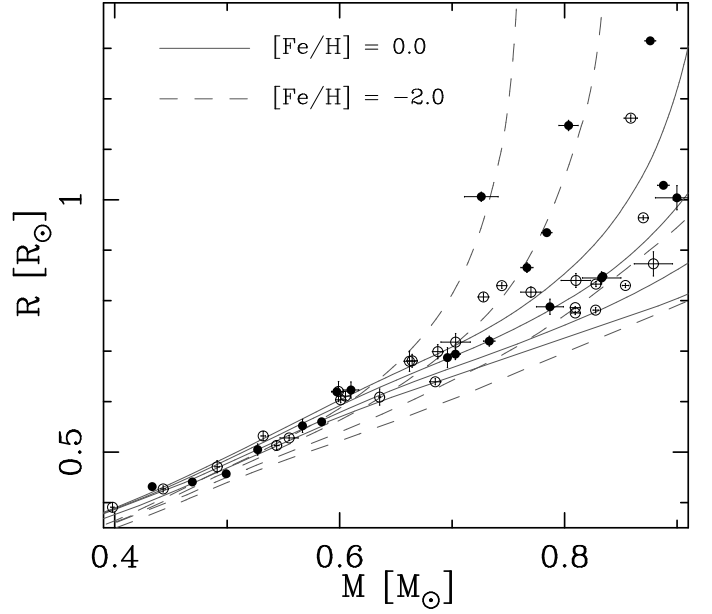


Fig. 11. Mass-radius diagram for well-measured DEBs with at least one component in the mass range of $0.4-0.9 M_{\odot}$. Solid symbols represent primary (more massive) and open symbols secondary components. Dartmouth isochrones of solar composition (solid) and metal-depletion (dashed) are plotted for ages 1, 5, 10 and 14 Gyr.

tionary relation. Usually the secondary appears to be oversized (over-aged), but literature examples show that it may also be the primary.

According to the recently popular explanation of the observed oversized radii of low-mass DEBs, the fast rotational velocity, resulting from short orbital period, is responsible for higher level of activity, which inhibits the effectiveness of convection, leading to larger radii and lower temperatures (Chabrier et al. 2007). Long-period systems should thus be less active and have their radii and temperatures better reproduced than the short-period binaries. One can see in Fig. 10 that AK For does not follow this trend, despite being active. On the other hand, the $0.85 M_{\odot}$ secondary of a 4.8 d system V636 Cen is both active and inflated (Clausen et al. 2009), as well as both components of the 41 d period LSPM J1112+7626 (Irwin et al. 2011). It is worth mentioning that the Ca II emission lines are much stronger in AK For than in V636 Cen, which rotates a bit faster. Some of the components of longer-period systems found in globular clusters, like the secondaries of V66 and V69 in M4 (Kałuzny et al. 2013) or the primary of V54 in M55 (Kałuzny et al. 2014), also seem to be oversized, but their activity is not seen. In the same time we also know short-period active systems, like ASAS J045304-0700.4 ($P = 1.77$ d; Hełminiak & Konacki 2011), MG1-78457 ($P = 1.59$ d), MG1-506664 ($P = 1.55$ d; Kraus et al. 2011), or KOI-126 BC ($P = 1.77$ d; Feiden et al. 2011), where the model radii fit the observed ones. In the light of recent findings it is difficult to draw any conclusions about the connection between rotation, activity, radii and effective temperatures.

5. Summary

We present the first full orbital and physical analysis of the brightest low-mass detached eclipsing binary known to date – AK Fornacis. The system's characteristics (brightness, long orbital period) allow for a very accurate determination of masses, and fairly good of the radii. The precision in masses is very high

but still somewhat limited by the stellar activity. The object's characteristics are unusual among the known low-mass DEBs, as despite being very active, properties of both components are nicely reproduced and the age can be determined fairly well as for a main sequence system. The reason for the enhanced activity can not be the rotation, as AK For rotates rather slowly, although faster than typical, single field K and M dwarfs.

The major observational disadvantage of the system is its period, very close to 4 d, which makes the eclipses not always possible to observe from one place during a single season. Further work on this system requires continuous photometry from a global network of telescopes, such as WET (Nather et al. 1990), *Solaris* (Konacki et al. 2012) or HAT-S (Bakos et al. 2009), in order to derive some of the crucial parameters, like the radii or temperatures, even more accurately.

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